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constant in the region where $M_1=\max_{[|\phi_1]}|\Delta\phi|$. One example of such cases
are axion-like fields, where $k(p)$ is constant up to the point where the axion decay constant
becomes of the order of the Hubble scale, $f_a\sim H$. In the case of an axion potential of the
form $V(\phi) = m^2\phi^2/2$, the bounce can be found in [@Hwang:2005he], where a
realistic inflation model was constructed. Conclusions =========== We have shown that a
model with the potential and the effective potential $V(\phi)+\Delta V(\phi)$ can be employed
to describe the effective physics between the quantum state of the field at a certain point of

time and the resulting classical background geometry. In particular, we have constructed a model for the reheating of the universe after inflation, which is driven by a phantom-like scalar field. The existence of a minimum of the effective potential is needed for the phenomenon to be possible, however we do not require the phantom field to be globally unstable to be able to obtain a stable bounce. The large field inflationary model employed in the following analysis can be implemented in a (nearly) realistic extension of the standard model. Namely, if we assume that the observed tensor-to-scalar ratio r is due to the value of the Hubble scale during inflation, it is natural to suppose that the tensor-to-scalar ratio is of the order of 0.1, which is also the order of magnitude of the critical energy density at the end of inflation. To consider only the spatial homogeneous modes and neglect the back-reaction of the inhomogeneities, the scale factor of the universe can be approximated by a linear function in the conformal time η during inflation. 4bc0debe42

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